

High-resolution Fourier-transform ultraviolet-visible spectrometer for the measurement of atmospheric trace species: application to OH

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A compact, high-resolution Fourier-transform spectrometer for atmospheric near-ultraviolet spectroscopy has been installed at the Jet Propulsion Laboratory's Table Mountain Facility (34.4 °N, 117.7 °W, elevation 2290 m). This instrument is designed with an unapodized resolving power near 500,000 at 300 nm to provide high-resolution spectra from 290 to 675 nm for the quantification of column abundances of trace atmospheric species. The measurement technique used is spectral analysis of molecular absorptions of solar radiation. The instrument, accompanying systems designs, and results of the atmospheric hydroxyl column observations are described. © 2001 Optical Society of America

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1. Introduction

Long-term measurements of transient species in the HO_x, NO_x, ClO_x, and BrO_x families are important for the validation of models of stratospheric chemistry and transport and to verify that the phaseout of chlorofluorocarbons mandated by the Montreal Protocol is having the desired effect on stratospheric composition. The free radicals OH, NO₂, NO₃, and BrO are key intermediates in catalytic ozone destruction cycles and are suitable candidates for detection by ground-based UV-visible spectroscopy. Of these species, the OH radical plays the most important role from the surface through the mesosphere.

A number of techniques have been employed to characterize the distribution of upper-atmospheric OH including *in situ* laser-induced fluorescence,^{1,2} far-infrared spectroscopy,³⁻⁶ and near-UV spectro-

scopy.⁷⁻⁹ Only a few ground-based installations that use either a polyetalon Fabry-Perot,¹⁰⁻¹² a grating instrument,^{13,14} or a Fourier-transform spectrometer (FTS)¹⁵ have an extended data record for the vertical column of OH that would allow the analysis of temporal variations and their relation to atmospheric chemistry.

Infrared FTS's have been used effectively for many years in atmospheric composition measurements.^{5,16-19} Significant strides in the application of FTS's to the study of visible and near-UV spectra have been made relatively recently²⁰⁻²⁴; however, their use in atmospheric remote sensing has been limited.^{15,25,26} Although the FTS design suffers from a multiplex disadvantage²⁰ compared with a grating spectrometer with an array detector when used in the UV, the combined advantages of high throughput, high spectral resolution, broad spectral coverage, and wavelength calibration stability and accuracy,²⁷ all in a compact design, make the FTS an attractive alternative for atmospheric field measurements. For these reasons a Fourier-transform ultraviolet spectrometer (FTUVS) was developed for use in the visible and near UV to detect and quantify terrestrial atmospheric absorptions of solar radiation.

When the FTUVS is used in the direct solar imaging mode, the 27-day rotation period at the solar equator, which induces a Doppler shift (maximum 0.38 cm⁻¹) in the solar spectral features at opposing solar limbs, can be utilized to discriminate the ter-

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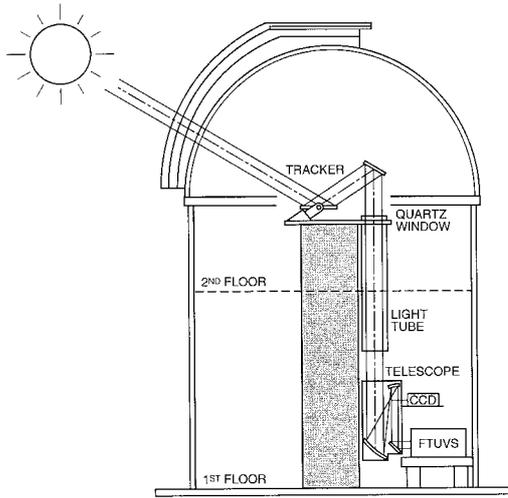


Fig. 1. FTUVS instrument system at the NASA JPL Table Mountain Facility.

restrial absorptions within the solar Fraunhofer structure.^{13,25} This mode places special requirements on the object tracking and imaging system that delivers light to the spectrometer.

The design of an instrument used to select and track on a solar limb, and achieve high spectral resolution to discriminate the terrestrial absorption features, is described in this paper. A demonstration of the instrument, applied to the measurement of terrestrial OH column abundances since July 1997, is also presented.

2. Instrument Description

The FTUVS system is located at the Jet Propulsion Laboratory's (JPL's) Table Mountain Facility (34.4 °N, 117.7 °W, elevation 2290 m). The instrument is comprised of the three basic subsystems as shown in Fig. 1: (1) a heliostat for tracking the Sun or Moon; (2) a beam-defining telescope with aperture and field stops, which also provides a feedback image for closed-loop tracking; and (3) an interferometer that records spectra of the selected source (Sun, Moon, or zenith sky). In addition, a data-acquisition and instrument control system allows collection and coadding of multiple scans to improve the signal-to-noise ratio (SNR).

A. Heliostat

The heliostat, constructed by DFM Engineering of Longmont, Colo., is located on the second floor of the building within the dome structure and is exposed to ambient conditions. The heliostat primary and secondary mirrors, 50.8- and 32-cm Zerodur substrates, respectively, both aluminum coated with SiO₂ and MgF₂ overcoats, were sized to provide an unvignetted 0.5-deg field of view and a 20-cm-diameter beam. The equatorially mounted heliostat primary mirror is driven in right ascension and declination to track any celestial object to within an accuracy of 1.8 arc sec/min. The light beam from the heliostat passes

through a 25-cm-diameter UV-grade fused silica before reaching the first-floor temperature-controlled instrument room.

B. Telescope

In the instrument room the beam passes through the telescope aperture stop before reaching the afocal Gregorian telescope primary. The telescope consists of confocal parabolic mirrors coupled to an Invar metering rod for thermal stability. The telescope off-axis parabola design is free from spherical aberration, coma, and astigmatism. The design also minimizes reflection losses by use of only two surfaces to compress the beam. The off-axis parabola telescope optics are aluminum coated, SiO₂ and MgF₂ overcoated, with a $\lambda/8$ wave front.

The telescope primary is an $f/3.2$, 20-cm-diameter off-axis segment of a parabolic mirror that focuses the light beam onto a reflective, oxygen-free copper field stop. The field stop aperture allows part of the focused image to pass through to one of two telescope secondary mirrors where it is recollimated by an $f/3.2$, 6.4-cm-diameter parabolic mirror for delivery to the FTUVS interferometer. The rest of the image is reflected by the aluminum-coated field stop toward a 3.3-cm-diameter $f = 30.5$ -cm parabolic secondary. This mirror collimates the light and delivers it through neutral-density filters to a focusing lens and a CCD camera. The image is displayed on a television monitor, captured by a frame grabber board, and processed with image analysis software developed at JPL. The image position in the CCD frame is used to provide closed-loop tracking correction signals to the heliostat motor driver.

C. Interferometer

The collimated light beam from the telescope is directed toward the FTUVS interferometer (shown schematically in Fig. 2). The FTUVS was designed and built at the JPL. It is a flex-pivot-supported, voice-coil-actuated, active-aligned plane mirror interferometer with an etendue of 2×10^{-4} cm² sr.²⁸ Plane mirrors were chosen to minimize the number of reflections and reduce wave-front distortion. The interferometer moving mirror mechanism is a parallelogram consisting of four tubes constructed from 152- μ m-thick rolled Invar for thermal and mechanical stability, connected by eight Bendix flex pivots. The lower, horizontal side of the parallelogram supports a mirror. This arm of the parallelogram is driven by a voice coil actuator along an essentially frictionless porch swing path moving the mirror horizontally 5 cm.

The mirror is moved at a constant velocity ($\pm 0.15\%$ rms) by a Hewlett-Packard (HP) Model 5507A laser positioning system servo-axis driver. This system uses a 2.1-MHz Zeeman split, linearly polarized, two-frequency He-Ne laser. The polarized beams are split one each into the two interferometer arms by a small polarizing spot coating on the beam-splitter front surface. The Doppler-shifted frequency of the beam from the moving mirror arm is compared with

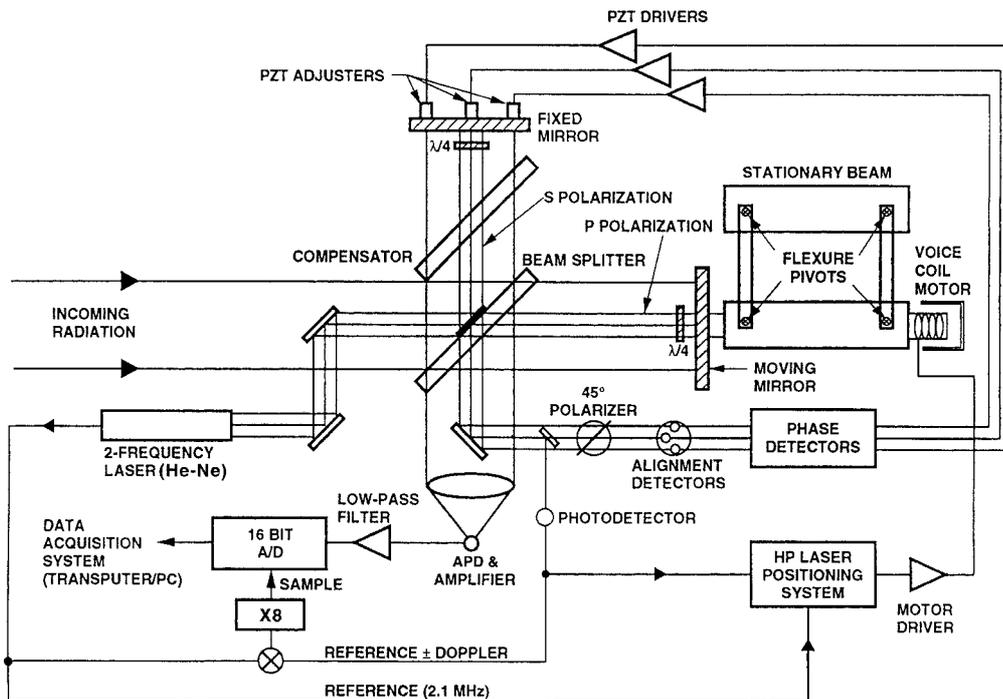


Fig. 2. FTUVS interferometer subsystem schematic. PZT, piezoelectric transducer; A/D, analog to digital converter; APD, avalanche photodiode.

the laser reference frequency for measurement and control of the scan mirror velocity. We also use the He-Ne beam to produce an interferogram sampling trigger signal by mixing the 2.1-MHz Zeeman split laser reference frequency with the Doppler-shifted moving mirror arm laser frequency to produce a laser fringe marker.

The intrinsic mechanical alignment accuracy of this system is better than $1 \mu\text{rad}$ over 1 cm of travel. To achieve the desired mirror tilt stability of $0.1 \mu\text{rad}$ over 5 cm of mechanical travel (to maintain high fringe contrast at high spectral resolution), an active tilt compensation system was added by use of a servo-controlled, piezoelectric actuator driving the fixed mirror to match the moving mirror tilt.²⁹ The servo is activated by three miniature p-i-n photodiodes sensing the phase of the 6-mm-diameter HP laser positioning system He-Ne laser reference beam at three places in the beam. A phase comparator determines the phase mismatch and produces the voltage that is required to drive the three piezostacks behind the fixed mirror to null the phase mismatch. The advantage of a heterodyne (Zeeman) laser system over a dc system employed in some commercial FTS designs is that alignment lock can be achieved and maintained without mechanical dithering, which is required to produce fringes.

Optics for the interferometer, beam splitter, compensator, and mirror substrates were figured by IC Optical of Great Britain. Coatings on the interferometer mirrors are aluminum with a SiO_2 over-

coat. The interferometer beam splitter has a multilayer dielectric coating, applied by Barr Associates, Inc., with 30–50% reflectivity and 250–650 nm at a 45-deg incidence angle. The beam splitter also has a central spot coating that divides the two orthogonally polarized components of the He-Ne reference laser into the two arms of the interferometer.

For measurements of the OH column abundance, a narrow-bandpass interference filter, 75 mm in diameter, is placed at the entrance to the interferometer. Its bandpass is centered at 307.8 nm where the peak transmission is 45%. The full-width at half-maximum (FWHM) of this filter is 3.8 nm, and out-of-band rejection is greater than 10^{-7} . The filter is used to isolate the strongest OH rotational absorption lines in the (0,0) vibrational band, $X^2\Pi \rightarrow A^2\Sigma$ electronic transition.

The detector is a 5-mm-diameter UV-enhanced avalanche photodiode with a quantum efficiency of greater than 6% in the UV and a current gain of 150. The FTUVS instrument system has been operated successfully from 290 to 675 nm.

D. Data System

In digitizing the signal, it is necessary to subdivide the laser fringes because the OH lines are at a wave number close to a multiple of the laser wave number. It is therefore not possible to use an electronic bandpass filter to isolate one laser frequency alias (of a width of 7899 cm^{-1}) and sample it simply with the laser fringes. To accomplish the subdivision of laser

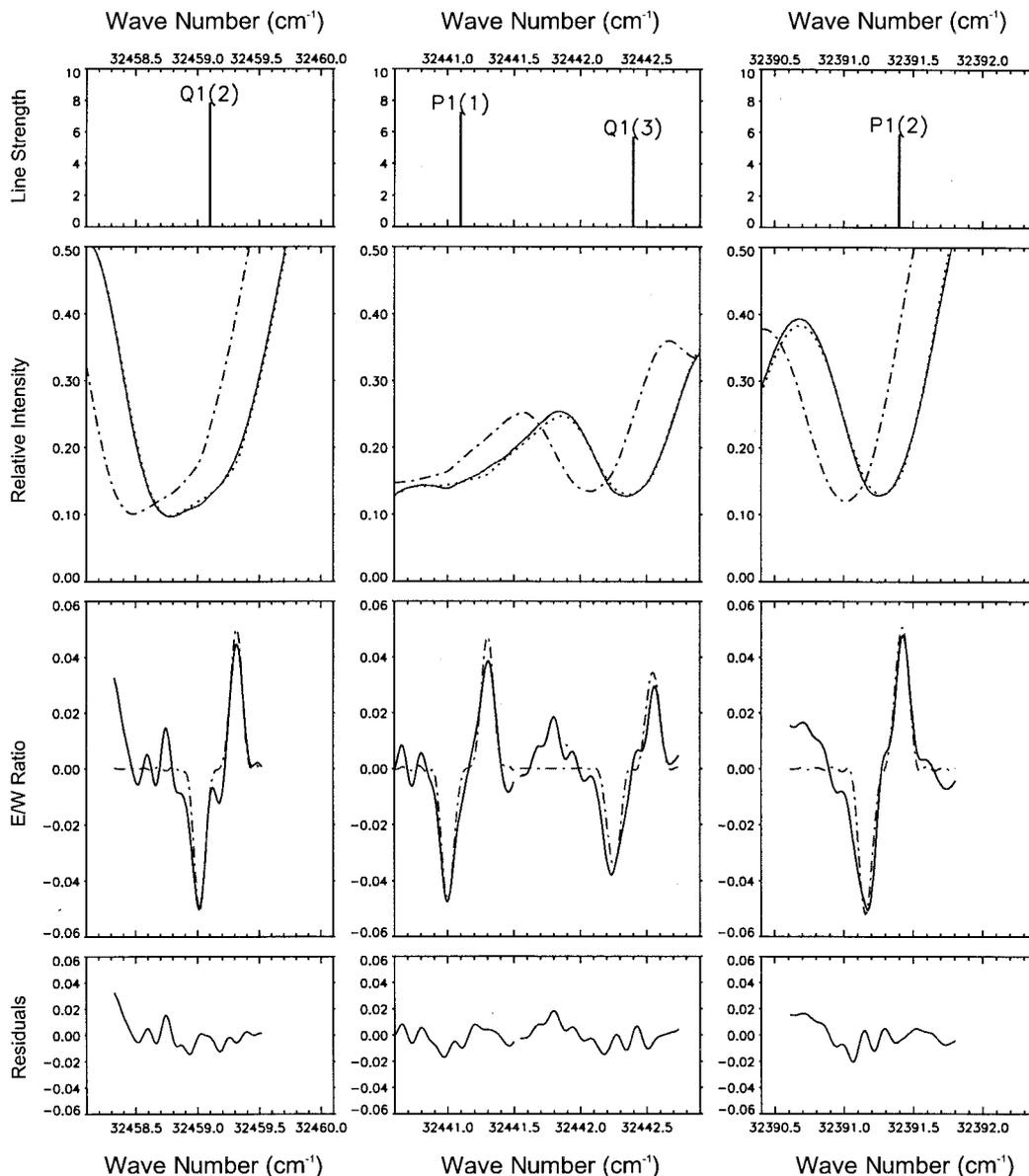


Fig. 3. Spectral retrieval of terrestrial OH lines by use of solar Doppler differencing: top row, positions and relative intensities of strong OH lines; second row, west solar limb spectrum (dot-dash curve), east spectrum (solid curve), and shifted west spectrum (dotted curve); third row, ratio spectra (solid curve) and fitted OH reference spectra (dot-dash curve); bottom row, residuals after fit of observed lines.

fringes, we used an analog phase-locked loop. The voltage-controlled oscillator of the phase-locked loop was selected for its low jitter noise. The laser fringe phase detector has two symmetrical outputs and was implemented in a fast integrated circuit to react to errors of the order of a few nanoseconds. A 125-Hz active filter with differential inputs is used to reject the frequency of the laser zero crossings, which is typically in the 6–12-kHz range. In addition, the offset of this active filter was nulled to avoid a static phase error when the circuit is fed with a constant frequency. This reduces phase detector outputs in the constant velocity mode of operation. The resulting digitized raw interferogram data are recorded to a large hard disk. The system can sustain data rates up to 10^5 samples/s, at 2 bytes/sample.

E. Data Reduction

The instrument produces nearly single-sided interferograms that are individually phase corrected³⁰ and Fourier transformed. These spectra are then averaged for the desired period of measurement. For the OH column observations, spectra are collected over a 15-min period at each solar limb. These summed east and west limb solar spectra are interpolated on a finer grid when the spectrum is convolved with a sinc function of a width determined by the instrument scan parameters. The necessary spectral shift (approximately 0.28 cm^{-1}) to be applied to the west limb data to match the solar Fraunhofer features of the east limb spectrum is determined by minimization of the residuals in a linear- least-

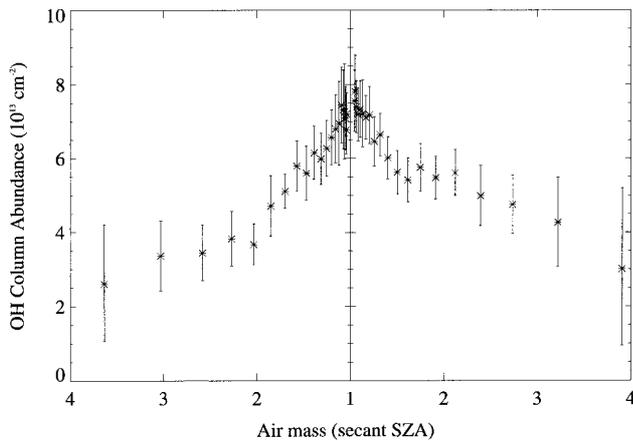


Fig. 4. Diurnal OH column abundance versus air mass measured on 11 May 2000, plotted out to four air masses. Vertical bars represent the $\pm 1\sigma$ statistical error in the fitted OH column. Points are plotted at the average value in the range of air masses during a 15-min integration period. This range increases with increasing air mass. SZA, solar zenith angle.

squares matching of the two spectra (see Fig. 3). The east limb spectrum is then divided by the shifted west limb spectrum. The terrestrial OH absorption feature appears in the ratio as either a reduction in the ratio (terrestrial OH absorption in the east limb spectrum divided by the west limb solar background) or as an increase in the ratio (east limb background divided by OH absorption in the west limb spectrum). The residual baseline slope is then removed, and the absorption signatures of five to seven OH absorption lines are fitted to a synthetic Doppler line shape for OH at 250 K, convolved with the instrument line shape. This average OH column temperature of 250 K was chosen based on a calculated OH density-weighted model of the temperature profile.³¹ Temperature sensitivity studies of the integrated OH line absorption cross sections^{32,33} demonstrated that the retrieved column abundance varied only $\pm 7\%$ for a ± 30 K temperature change of approximately 250 K. The synthetic OH absorption line fit is adjusted to minimize the square deviation between the fit and the measured OH feature. The resulting measured line intensity and the calculated absorption cross section of the line are then used to determine the slant column of OH. We convert this to vertical column abundance by dividing the slant column by the air mass (approximately the secant of the solar zenith angle). Observations are obtained for air masses up to five (object zenith angles of 78 deg). Below this air mass, refraction has an insignificant effect on the slant column path length.

3. Results

The derived OH column is obtained from a weighted average, proportional to the line strength, of the strongest OH absorption lines. Over 180 OH diurnal column data sets were acquired beginning in July 1997 for solar zenith angles of 10–78 deg (air masses of 1.02–5.0). Figure 4 presents a single day of the

measured OH vertical column and its diurnal variation (local noon is near the center of the figure). Each point is derived from sequential 15-min integrations near the solar west and east limbs. Forty spectra are acquired, summed, and averaged to create each 15-min integrated data set. The SNR at local noon in these integrations, as measured by the $\pm 1\sigma$ noise level outside the bandpass of the optical filter, is approximately 1700:1 under clear-sky conditions (minimal source noise contribution from haze or clouds). For the case of FTUVS, where the noise contribution from the detector and from instrument emission is negligible, the theoretical SNR is given by³⁴

$$\text{SNR} = \delta\bar{\nu}\Phi\left(\frac{Nt}{\Delta\bar{\nu}}\right)^{1/2},$$

where $\delta\bar{\nu}$ is the spectral resolution of 0.06 cm^{-1} , Φ is the modulation efficiency of 0.6, N is the number of signal electrons of $1.5 \times 10^9\text{ s}^{-1}(\text{cm}^{-1})^{-1}$, t is the total integration time per interferogram of 22 s, and $\Delta\bar{\nu}$ is the width of the band-limiting filter of 400 cm^{-1} .

The value computed by the above expression yields a SNR of 1752, which agrees closely with the measured value.

The residuals from fits to the east/west ratio spectra exhibit features that are larger than the measured photon noise (Fig. 4). These features have also been observed in previous OH column abundance studies that used both grating^{13,14} and FTS's.^{15,25} The observed features could be attributable to spatial inhomogeneity in the solar emission spectra from the east and west solar limbs. Alternatively, the features could arise from time-dependent artifacts from the interferometer itself. Artifacts can result from coherent phase errors that occur at optical path differences that are not captured by the phase correction algorithm. These would include imperfect operation of the active alignment system, the shear motion combined with imperfect flatness of the moving mirror, and density fluctuations in the instrument associated with air currents.

In Fig. 4, vertical bars represent the change in the calculated OH column abundance that is produced by a 5% change in the χ^2 minimization when the synthetic spectra are matched to measured OH features. This is an indication of the measurement precision. For the data collected from November 1998 to June 2000, the median uncertainty on the fit by use of the Q1(2) line was $\pm 18\%$, and the median uncertainty on the fit by use of the P1(1) line was $\pm 15\%$.

We can obtain an upper limit on the uncertainty for the accuracy of the measurement by considering measurements collected on subsequent days at nearly the same solar zenith angle. The median standard deviation scatter of measurements at nearly the same solar zenith angle is $\pm 15\%$ when the column abundances collected from November 1998 to June 2000 for the strongest absorption lines are averaged. The range of the standard deviation is 10–22%, depending on the solar zenith angle. A more rigorous anal-

Table 1. Estimate of FTUVS Measurement Uncertainty

Error Sources	Estimate (%)	Description
Systematic		
Line center cross section	± 10	OH vibrational level measured lifetime uncertainty.
	± 7	Stratospheric temperature uncertainty in column.
Other absorptions	?	Possible unidentified absorption features at the location of OH lines.
Random		
Spectral distortions	± 10	Instrument stability, air currents, acoustic and electronic noise.
Photon noise	± 2	Poisson counting statistics.
Pointing system	± 10	Positioning drift on Sun. Spectrum Doppler shift changes. Integrated solar spectrum shape affected.
Constraining the fit (least-squares χ^2 optimization)	± 15	Match observed spectra with modeled line shape. Residuals assessed with a 5% variation about the χ^2 minimum.
Total	± 22	Accuracy

ysis of the entire data set, observation precision, and comparison with the modeled OH profile is in preparation.³⁵ The sources of these uncertainties, both systematic and random, are summarized in Table 1.

4. Summary

We have demonstrated a compact, high-resolution, actively aligned FTS for UV–visible atmospheric spectroscopy and have applied it to the measurement of OH column abundances. Design of the FTUVS, with emphasis on high throughput and the retrieval of multiple OH spectral lines, has resulted in a data record with higher temporal resolution and precision than any obtained previously. The small footprint of the instrument and its all-flexure design are compatible with future space flight applications aimed at high-resolution UV–visible spectroscopy of Earth and planetary atmospheres. Measurements of the OH column are being used to validate modeled stratospheric chemistry, including seasonal variability of OH related to satellite and ground-based measurements of upper-atmospheric ozone and water vapor,³⁴ as well as the short time-scale effects of solar coupling, dynamics, and related chemistry.

Analysis of similar terrestrial atmospheric column data sets for NO_3 , NO_2 , and O_3 observed with the FTUVS instrument with the following instrument operation modes are under way: (1) full disk lunar observations at $15,106 \text{ cm}^{-1}$ (662 nm) for nighttime abundances of NO_3 , (2) zenith sky observations of multiply scattered sunlight for the detection of O_3 and NO_2 , and (3) direct solar observations for the detection of NO_2 .

The capabilities of this instrument will be exercised in an attempt to measure other atmospheric free radicals, including BrO, IO, and OCIO, exploiting the broad spectral coverage and high resolution of this instrument. The FTUVS is colocated at the JPL Table Mountain Facility with several other instruments to investigate changes in, and the interactions of, HO_x , NO_x , and ClO_x families with ozone. Results from this complementary set of measurements are being analyzed to enable critical comparisons be-

tween observed and modeled ozone depletion and stratospheric change.

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